

## Problem Set 6: Spherical Collapse in an $\Omega_m = 1$ Universe

Due Friday, November 30

The evolution of a spherically symmetric, overdense perturbation in a  $k = 0$ ,  $\Omega_m = 1$  universe can be solved analytically up to the point of collapse, which makes it a very useful model despite its geometrical idealization. The evolution can be viewed from a Newtonian point of view, but the relativistic derivation is more broadly applicable. The basic trick is to recognize that, as a consequence of Birkhoff's theorem (in a spherically symmetric universe, only the interior mass matters), the perturbation itself must follow the equations of a  $k = +1$  Friedmann universe.

The spherical collapse model is discussed in §8.2 of Padmanabhan's book and in a famous paper by Gunn & Gott (1972, *Astrophysical Journal*, vol. 176, p. 1). Note, however, that the spherical collapse model discussed in those places is slightly different from the one described below (see part 3).

### (1) Perturbation overdensity

Consider an  $\Omega_m = 1$  universe containing a single, spherically symmetric, homogeneous overdense region (the perturbation). The background universe is described by the Friedmann equation for  $k = 0$  and the perturbation by the Friedmann equation for  $k = +1$  (since its density exceeds the critical density). The equations for the evolution of the background universe  $a(t)$  and for the radius of the perturbed region  $r(t_p)$  can therefore be written in the parametric form (see course notes, §3, or Ryden §6.1):

$$a = \frac{1}{2} a_* \eta^2, \quad t = \frac{1}{6} \frac{a_*}{c} \eta^3$$
$$r = r_* (1 - \cos \eta_p), \quad t_p = \frac{r_*}{c} (\eta_p - \sin \eta_p).$$

Use the conditions that (a)  $t$  and  $t_p$  must be equal, and (b) the perturbation vanishes at high redshift ( $\rho_p \rightarrow \rho$  as  $\eta$  and  $\eta_p \rightarrow 0$ ) to derive an expression for the overdensity  $\rho_p/\rho$  as a function of  $\eta_p$ . You may find it useful to introduce the quantity  $q$  which is the radius the perturbed region would have had if it had expanded at the unperturbed rate (i.e., if had not been overdense).

### (2) The linear regime

Show that the density contrast

$$\frac{\delta\rho}{\rho} \equiv \frac{\rho_p - \rho}{\rho} \propto t^{2/3}$$

when  $\eta_p \ll 1$ .

Show that the dimensionless velocity perturbation for  $\eta_p \ll 1$  is

$$\delta_V \equiv \frac{v - Hr}{Hr} = -\frac{1}{3} \left( \frac{\delta\rho}{\rho} \right),$$

where  $v = \dot{r}$  is the perturbation's expansion velocity and  $H$  is the Hubble parameter of the background universe.

Hint: You need to keep doing Taylor expansions in  $\eta_p$  until the quantities you are trying to compute don't vanish. It will be worth it in the end.

### (3) Turnaround

Let  $\delta_i$  denote the density contrast  $\delta\rho/\rho$  at some early time  $t_i$  ( $\eta_p \ll 1$ ), and let  $r_i$  denote the perturbation radius at this time.

Show that the perturbation expands to a maximum or “turnaround” radius

$$r_{ta} = r_i \left( \frac{5}{3} \delta_i \right)^{-1} \quad \text{at time} \quad t_{ta} = t_i \left( \frac{3\pi}{4} \right) \left( \frac{5}{3} \delta_i \right)^{-3/2}.$$

Show that the overdensity at turnaround is  $\rho_p/\rho \approx 5.5$ .

[Extra credit: The expressions for turnaround radius and time differ from those of Gunn & Gott (1972) by the factors of 5/3. Why?]

### (4) Collapse and virialization

Show that the perturbation collapses to zero radius at  $t_c = 2t_{ta}$ .

Assuming that non-radial velocities are negligible at turnaround, argue that after the perturbation collapses and reaches virial equilibrium (potential energy =  $-2 \times$  kinetic energy) it will have a radius  $r_c \approx r_{ta}/2$ .

If virialization occurs instantaneously at  $t = t_c$ , what will the density of the virialized object be relative to the mean density of the universe at  $t = t_c$ ?

For this part of the problem, assume that the virialized object that forms after collapse is constant density (and spherical).

### (5) Isothermal halo

A more realistic (though still idealized) model is that the post-collapse dark matter halo is a singular isothermal sphere, with  $\rho(r) \propto r^{-2}$  and enclosed mass  $M(< r) \propto r$ . Show that in this case the *half-mass radius* of the virialized halo is  $r_{1/2} = \frac{5}{12} r_{ta}$ .

### (6) Applications

If all went well, then in part 4 you derived the well-known lore that spherical perturbation in an  $\Omega_m = 1$  has a mean overdensity  $M/(4\pi R^3/3) = 178\rho_{\text{crit}}$  after it collapses and virializes, where  $R$  is the “virial radius.” Since the details of the physical argument are rough, and there is some cosmology dependence in any case, it is common to refer to the virial radius as  $R_{200}$ , within which the mass is  $M_{200}$  and the mean overdensity is  $200\rho_{\text{crit}}$ .

With this definition, show that the virial radius of a dark matter halo of circular velocity  $V_{200} = (GM/R_{200})^{1/2}$  is

$$V_{200} = 10H_0 R_{200} = 1000h \left( \frac{R_{200}}{1 \text{ Mpc}} \right) \text{ km s}^{-1}$$

and that

$$M_{200} = \frac{V_{200}^3}{10GH_0} = 2.3 \times 10^{14} h^{-1} M_\odot \left( \frac{V_{200}}{1000 \text{ km s}^{-1}} \right)^3,$$

where  $h \equiv H_0/100 \text{ km s}^{-1} \text{ Mpc}^{-1}$ .

What are the virial radius and virial mass of the Milky Way’s dark matter halo?